

# Search for coincident neutrino and gravitational wave emission from binary black hole mergers in the accretion disk of active galactic nuclei

Giacomo Bruno, Gwenhaël de Wasseige, Romain Gorski, Mathieu Lamoureux, Chris Raab, Matthias Vereecken\*

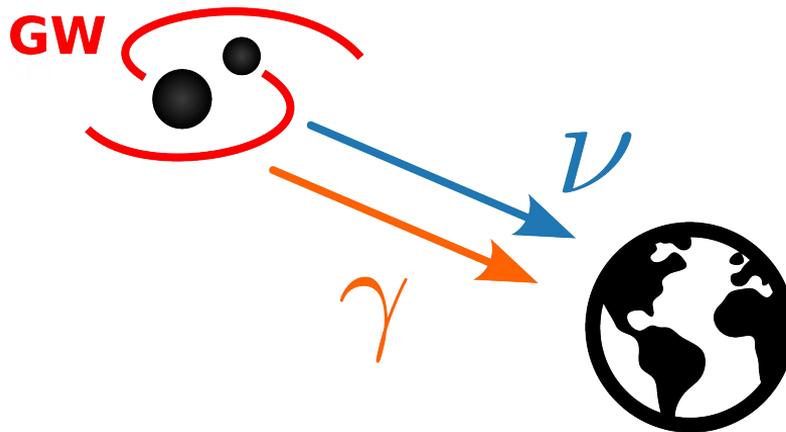
# Motivation

Multi-messenger astronomy

Several such searches already exist!

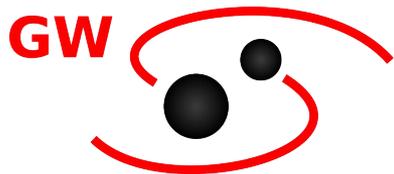
- LLAMA analysis (Bayesian)
- GW Follow-up searches by neutrino observatories (IceCube, KM3Net, SuperK,...)
  - ◆ Point-source searches at higher energies
  - ◆ Counting analysis at low energies

“General purpose” searches for common source of GW and  $\nu$



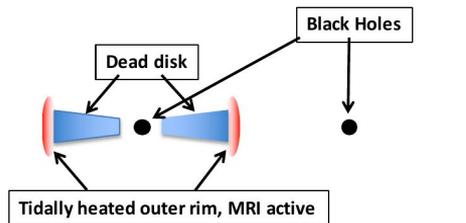
# How to get non-GW emission from BBH mergers?

Isolated binaries



✗ No emission

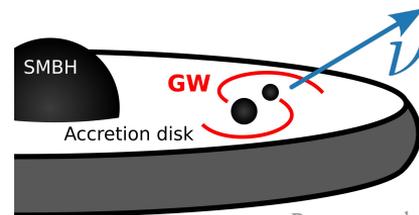
Dead disk



Perna et al. [1602.05140]

✓ Remnant disk reactivated

Binaries in AGN accretion disks



Bartos et al. [1602.03831]  
McKernan et al. [1907.03746]  
Kimura et al. [2103.02461]  
Tagawa et al. [2303.02172]  
...

- ✓ Many (heavy) black holes
- ✓ Frequent mergers
- ✓ Gas-rich environment

## Main take-away:

- *Steal as much as you can from GRB models*
- No clear singular model prediction

# BBH mergers in AGN

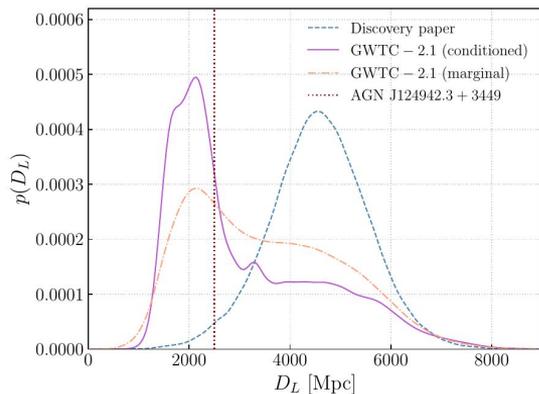
Localization alone can already probe this scenario

→ “Just” counting AGN

Bartos et al. [1701.02328]  
Veronesi et al. [2203.05907]  
Veronesi et al. [2306.09415]

GW190521 associated with ZTF19abanrhr from an AGN?

Graham et al. [2006.14122]

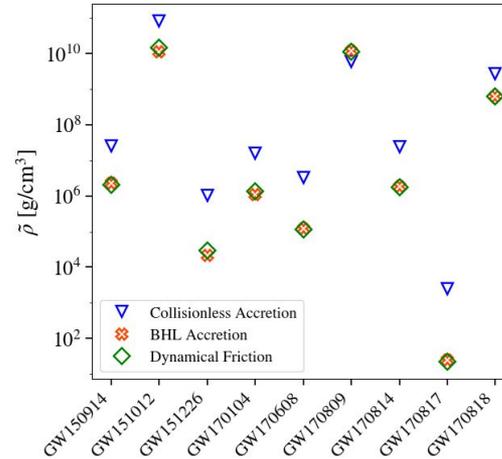
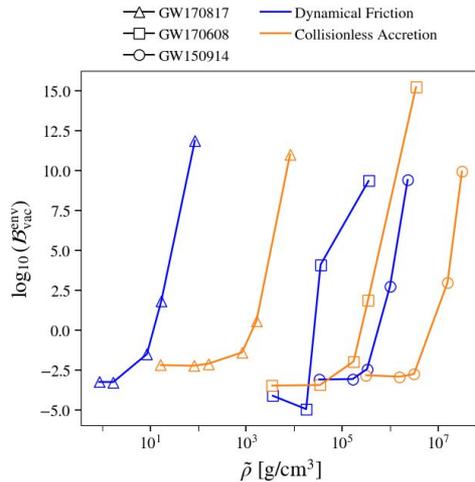


Morton et al. [2310.16025]

# Effects on waveform

Effect of gaseous disk environment on waveform

$$\rho \sim 10^{-8} - 0.1 (10^5 M_{\odot} / M_{\text{SMBH}})^{7/10} \text{g/cm}^3$$



Next generation will be able to probe this...

Santoro et al. [2309.05061]

# Joint $\nu$ -GW search for BBH in AGN disks

# Analysis

Given GW event: IceCube  $\nu$  in a time window T surrounding the event

Compare 2 hypotheses

$$\begin{array}{c} \mathbf{H}_0 \\ \nu \text{ sky} \\ = \\ \text{background} \end{array}$$

$$\begin{array}{c} \mathbf{H}_1 \\ \nu \text{ sky} \\ = \\ \text{background} \\ + \\ \text{point source} \\ @\text{AGN \& GW} \end{array}$$

# Neutrino point-source likelihood

$$\mathcal{L}_{S+B,\nu}(n_s; \hat{\mathbf{x}}) = \frac{e^{-(n_s+n_b)}(n_s+n_b)^{N_\nu}}{N_\nu!} \prod_{i=1}^{N_\nu} \frac{n_s \mathcal{S}(\hat{\mathbf{x}}_{\nu_i}; \hat{\mathbf{x}}) + n_b \mathcal{B}(\hat{\mathbf{x}}_{\nu_i})}{n_s + n_b}$$

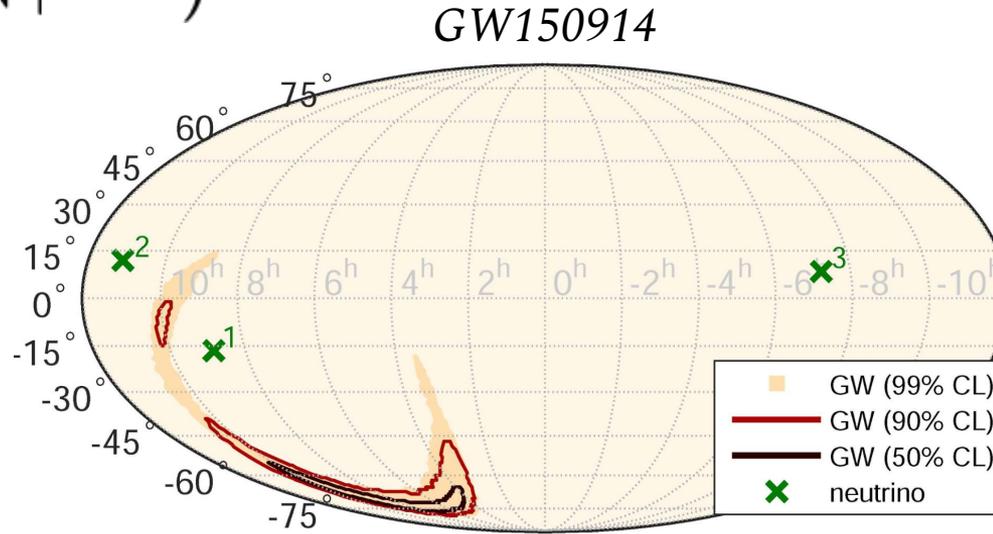
$\mathcal{S}(\hat{\mathbf{x}}_{\nu_i}; \hat{\mathbf{x}})$  Spatial likelihood signal neutrino

$\mathcal{B}(\hat{\mathbf{x}}_{\nu_i})$  Spatial likelihood background

Determine  $n_s$  using maximum-likelihood fit

# Direction prior (GW skymap)

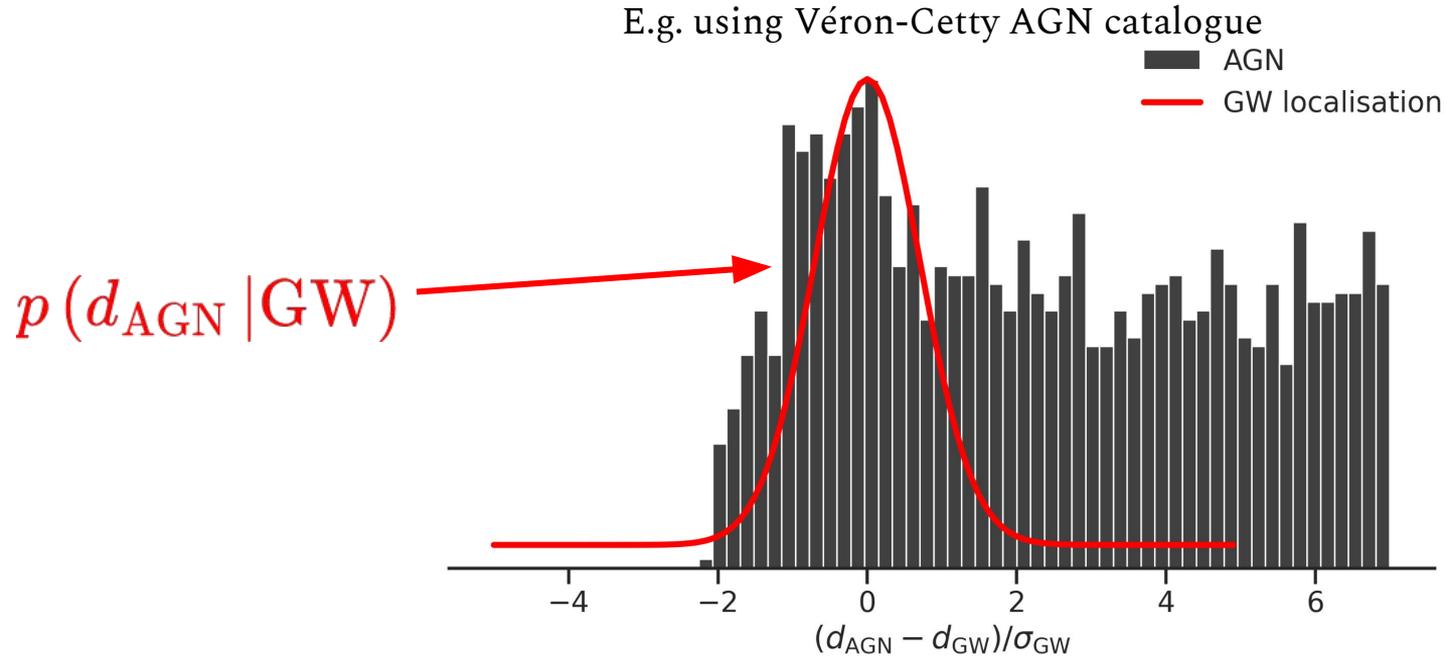
$$p(\hat{x}_{\text{AGN}} | \text{GW})$$



LIGO and Virgo Collab. [1602.03837]  
The IceCube Collab. [1602.05411]

# Distance prior

## Pixel-dependent GW distance probability



# Analysis

Best-fit = AGN which maximises TS

$$TS_{\text{AGN}} = 2 \ln \left( \frac{\mathcal{L}_{\nu, \text{S+B}}(n_s; \hat{x}_{\text{AGN}})}{\mathcal{L}_{\nu, \text{B}}} \times p(\hat{x}_{\text{AGN}} | \text{GW}) \times p(d_{\text{AGN}} | \text{GW}) \right)$$

Similar to existing  $\nu$ -GW unbinned ML TS

**New**

Significance:

- Perform trials with scrambled neutrino data
- Compare obtained TS-distribution with observed TS

# AGN catalog

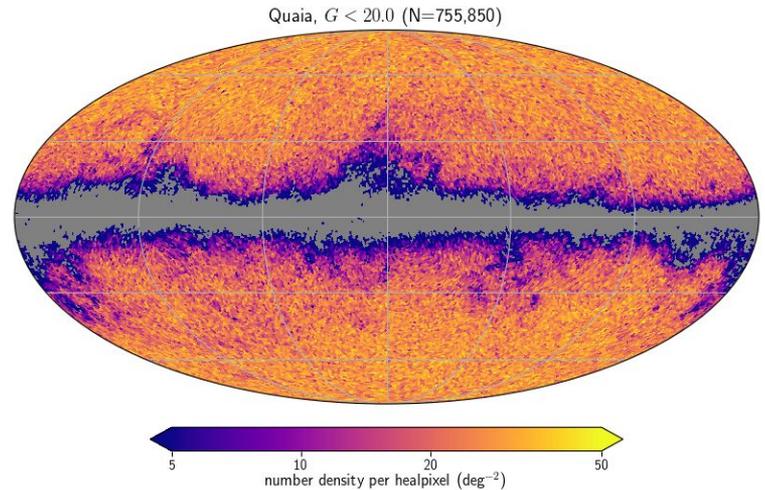
## Requirements

- Number of AGN
  - ◆ Too few AGN = miss coincidences
  - ◆ Too many AGN = lose sensitivity advantage
- Uniformity (cover full sky with similar depth)
- Completeness

# Quaia catalog

Quaia = Gaia DR3 quasar candidates (6.6M)  $\otimes$  unWISE (2B+)

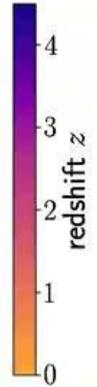
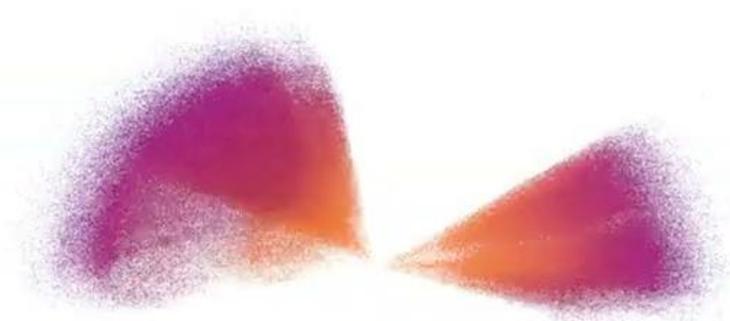
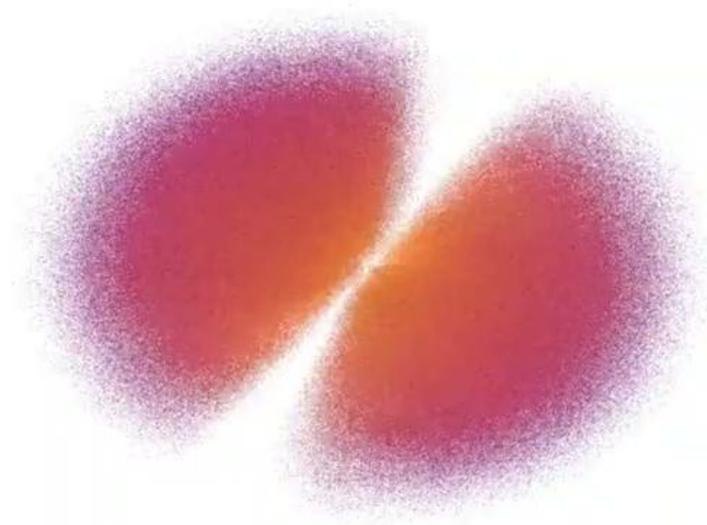
- 755 850 quasars with magnitude  $G < 20$
- 200 000 up to  $z \leq 1$
- $G < 19$ ,  $z \leq 1$  only 70k for improved purity



Storey-Fisher et al. [2306.17749]

Quiaia

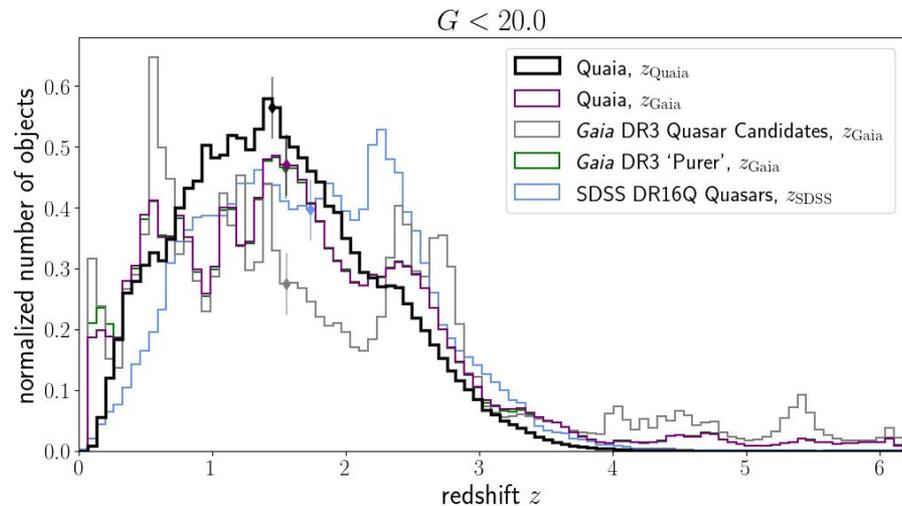
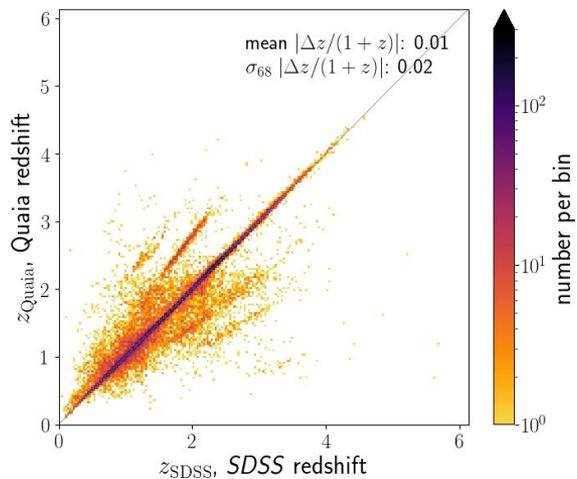
SDSS DR16Q



# AGN catalog - redshift information

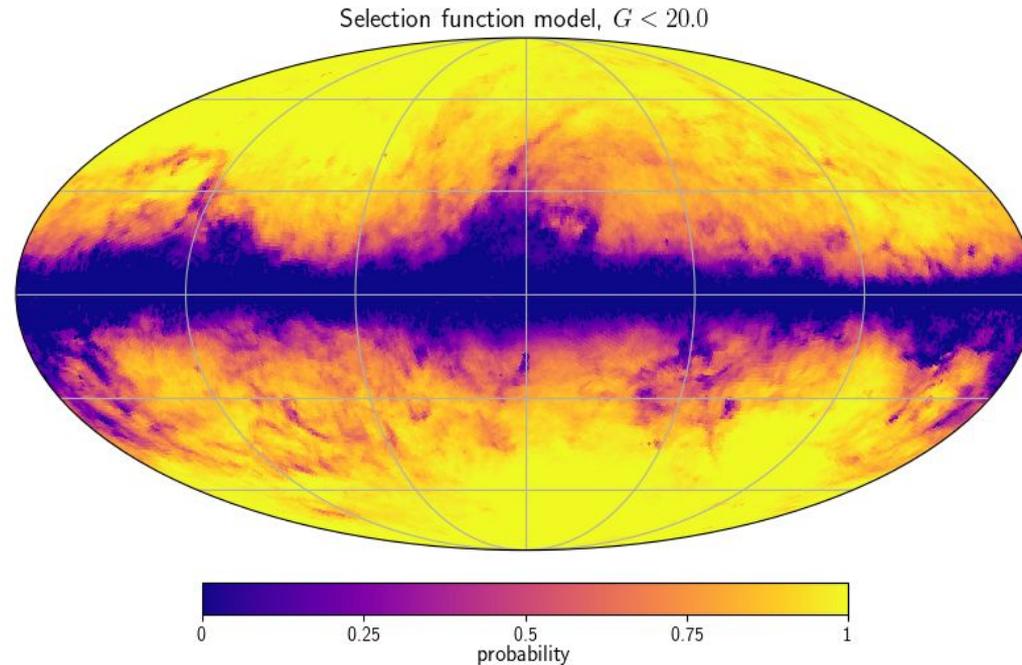
Photometric redshift, inherently less precise...

*But* error on redshift provided!



# AGN catalog - Selection function

Measure of completeness provided



# Next steps

## → Neutrinos

- ◆ Decide on time window (standard 1000 s)
- ◆ Start with high-energy neutrino sample (100 GeV - ~PeV)
- ◆ Add lower-energy samples in future:
  - GRECO (10 GeV - 100 GeV)
  - ELOWEN (1 GeV - 10 GeV)?
  - Hitspool (raw data, potentially even lower energy)?

## → Gravitational waves

- ◆ Add subthreshold GW - do we add new term to likelihood?

## → AGN Catalog

- ◆ Once analysis is designed, go back and evaluate properties of catalog